Based on [2307.05176]

Testing Sterile Neutrinos with Direct Detection and Spallation Source Experiments

DAVID ALONSO-GONZÁLEZ*

with D.W.P. Amaral, A. Bariego-Quintana, D. Cerdeño & M. de los Ríos

MultiDark20 (October 26th, Gandia)

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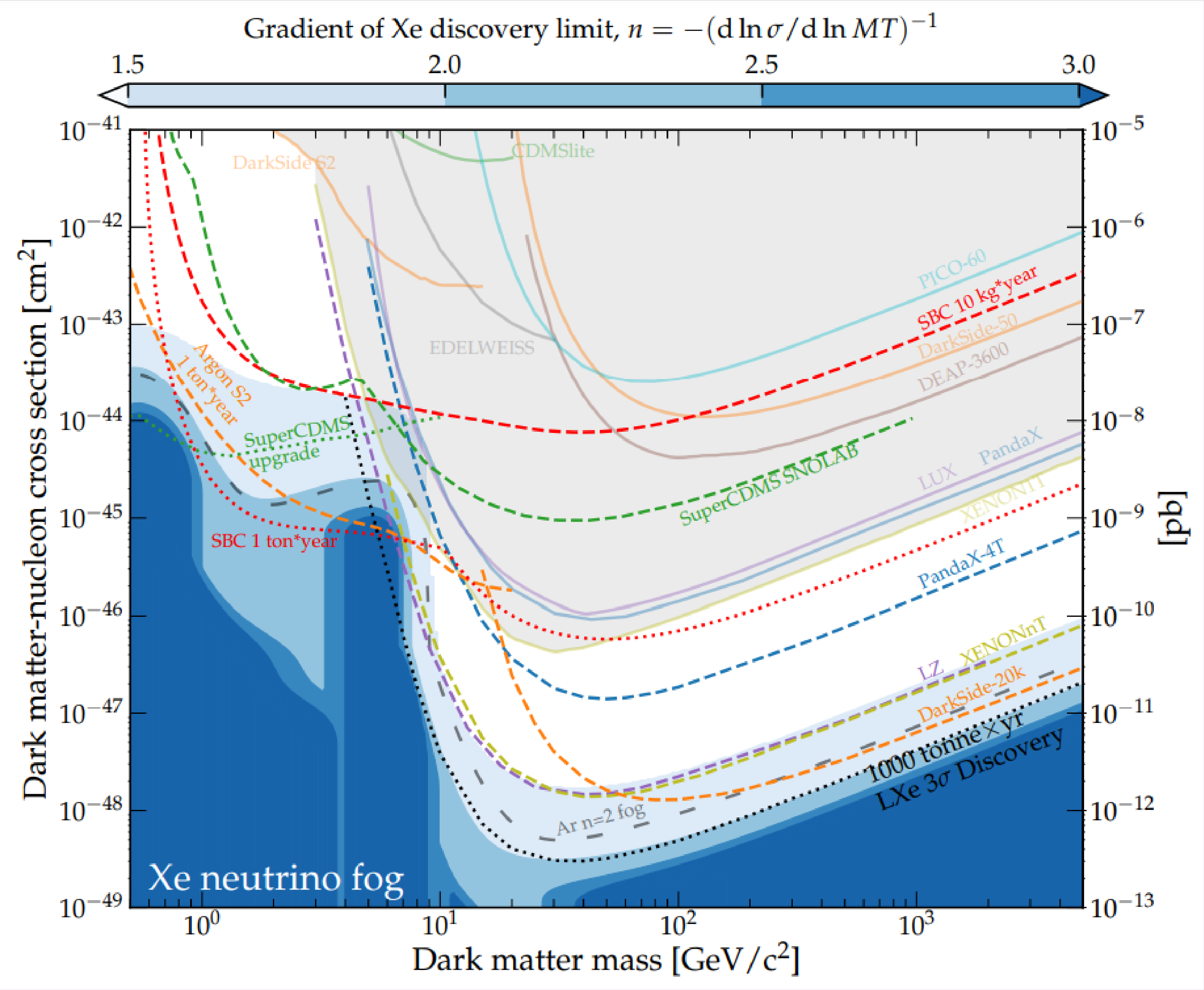

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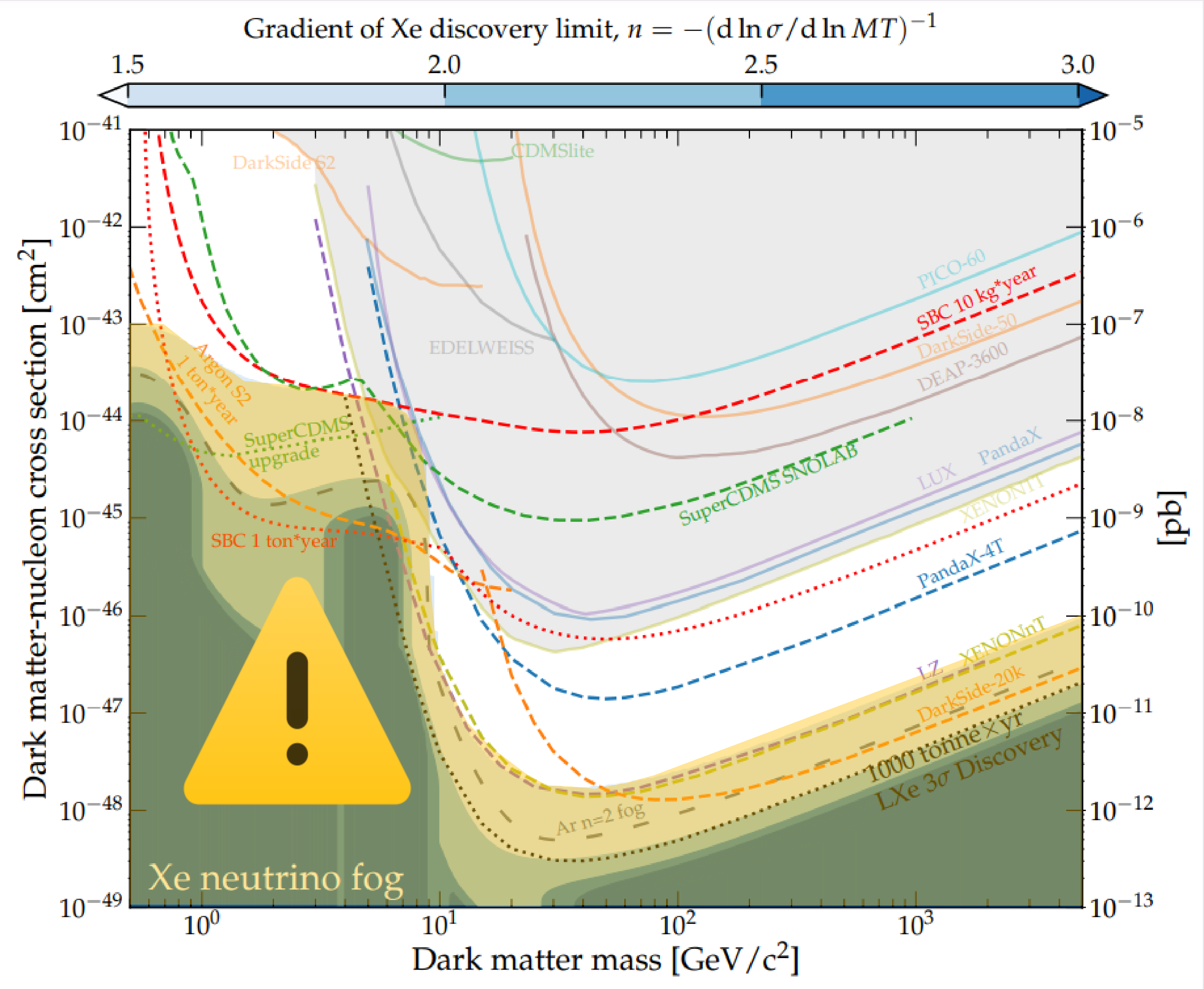

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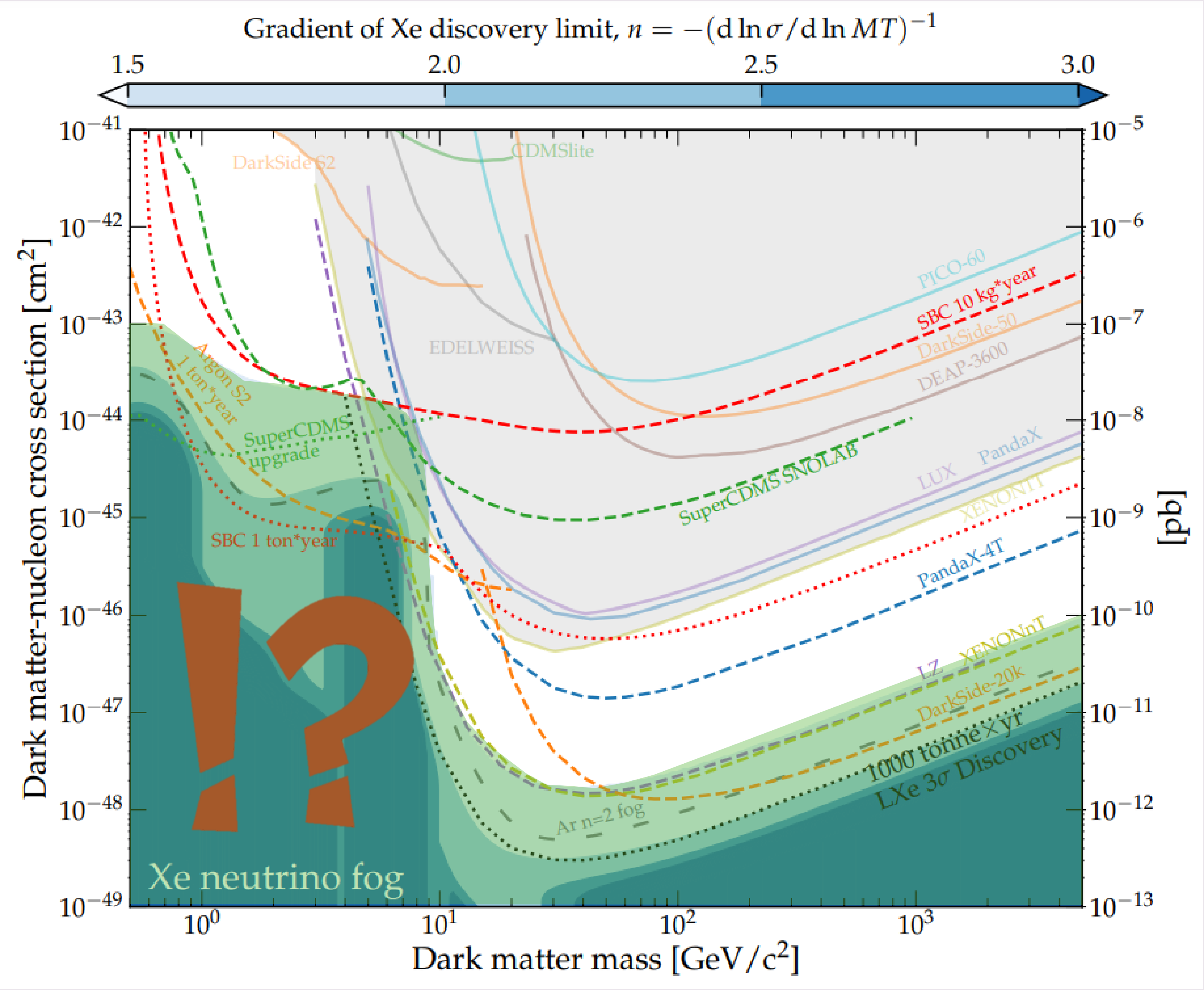

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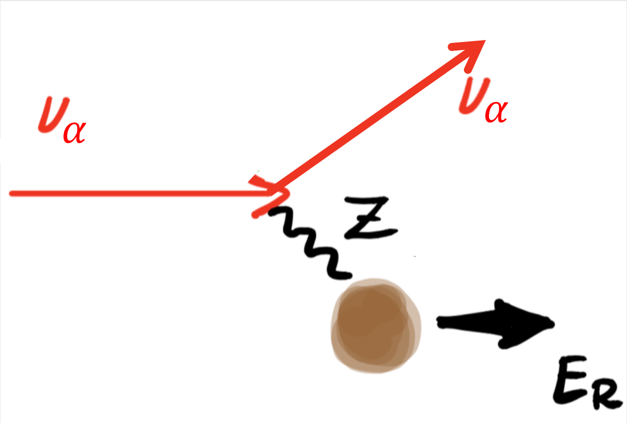
Coherent Elastic Neutrino – Nucleus Scattering

…or **CEνNS** for friends



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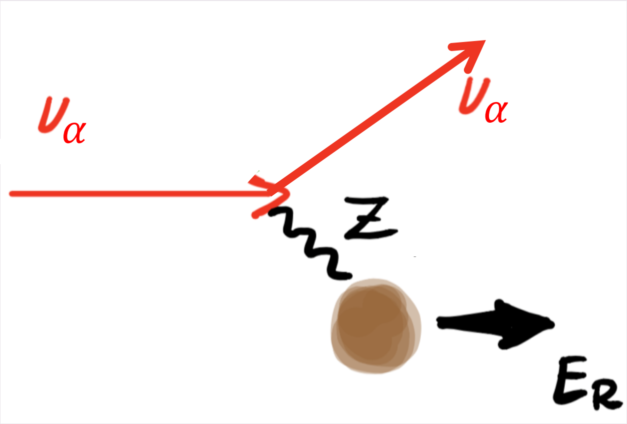

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[Freedman (1974)]

$$\frac{d\sigma_{\nu N}}{dE_R} = \frac{G_F^2}{4\pi} Q_v^2 m_N \left(1 - \frac{m_N E_R}{2E_\nu^2}\right) F^2(E_R)$$

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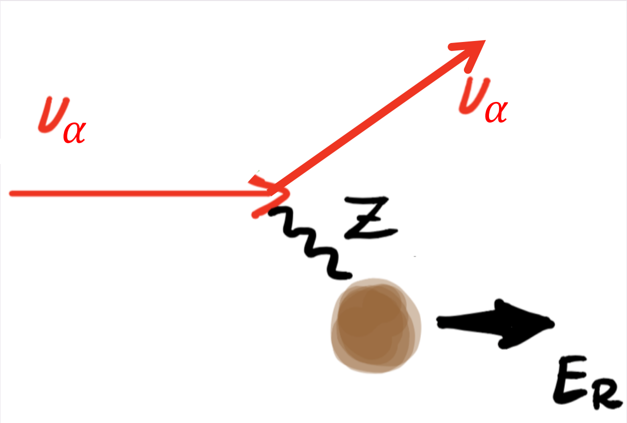

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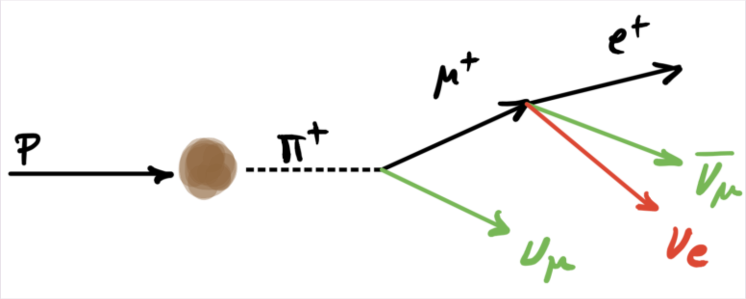
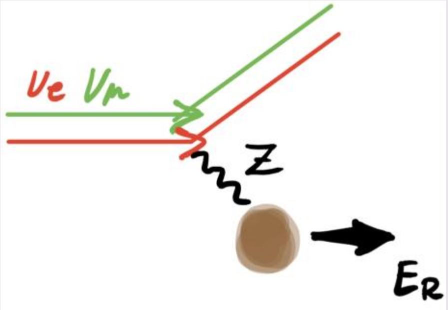
CEνNS at Spallation Sources

like **COHERENT**





$$N_{\mathrm{CE}\nu\mathrm{NS}} = \sum_{\nu_\alpha} N_{\mathrm{targ}} \int_{E_{\mathrm{th}}}^{E_R^{\mathrm{max}}} \int_{E_\nu^{\mathrm{min}}}^{E_\nu^{\mathrm{max}}} \frac{\mathrm{d}N_{\nu_\alpha}}{\mathrm{d}E_\nu} \epsilon(E_R) \frac{\mathrm{d}\sigma_{\nu_\alpha N}}{\mathrm{d}E_R} \mathrm{d}E_\nu \, \mathrm{d}E_R$$

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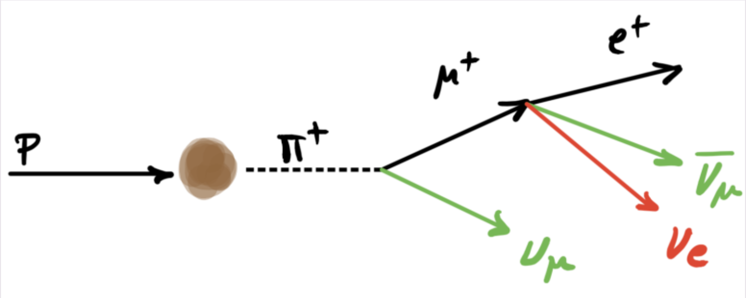

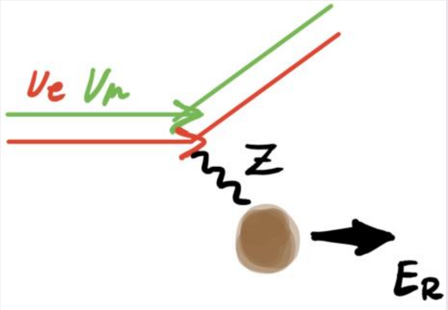

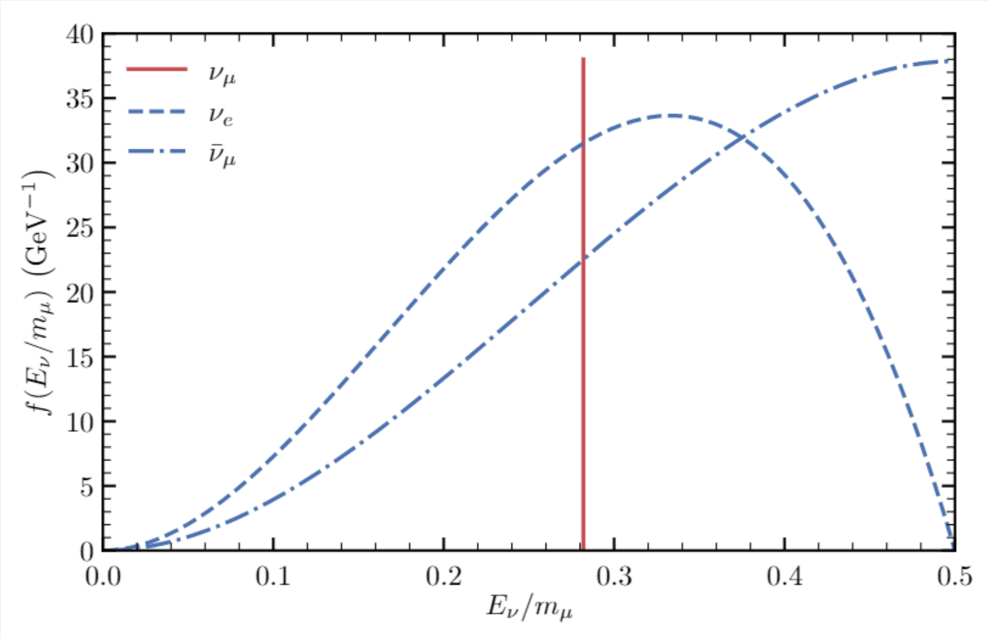

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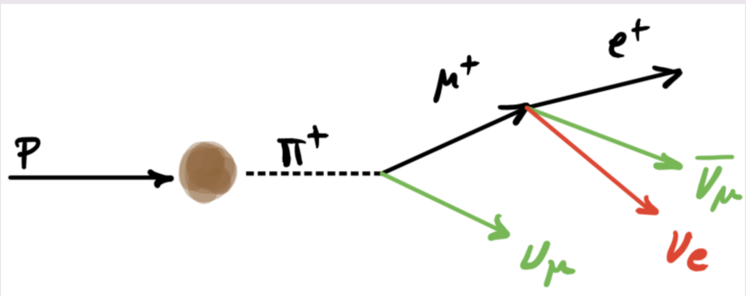

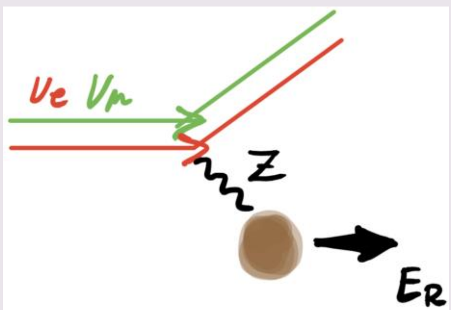

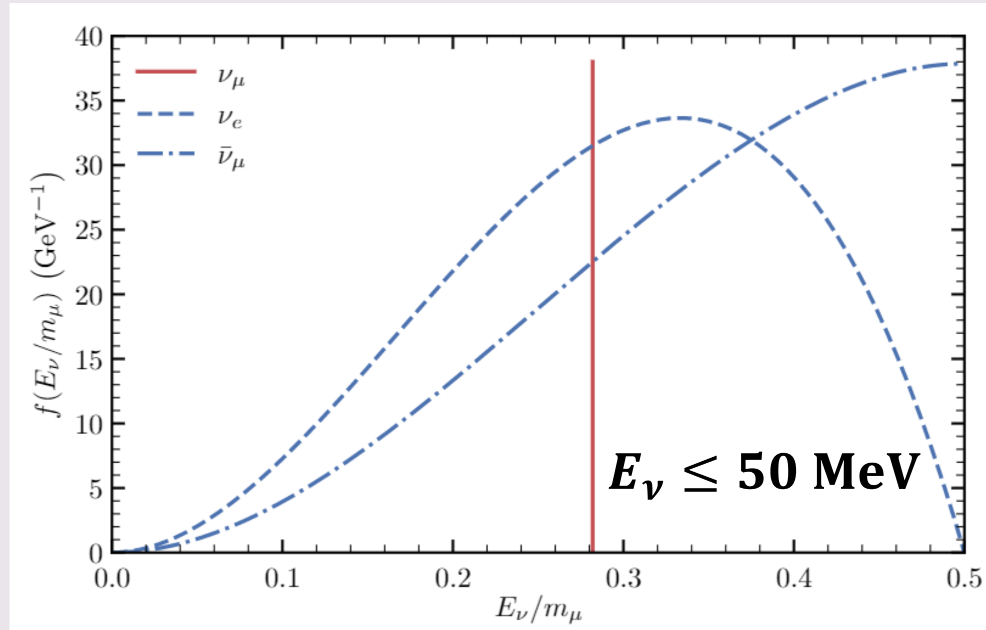

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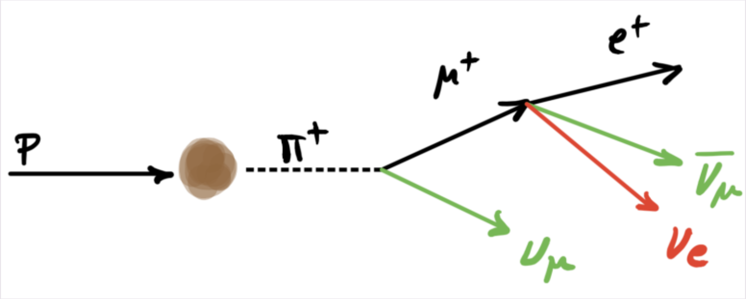
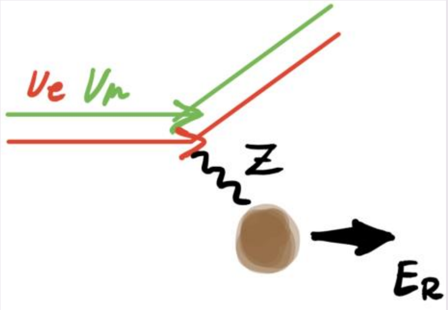
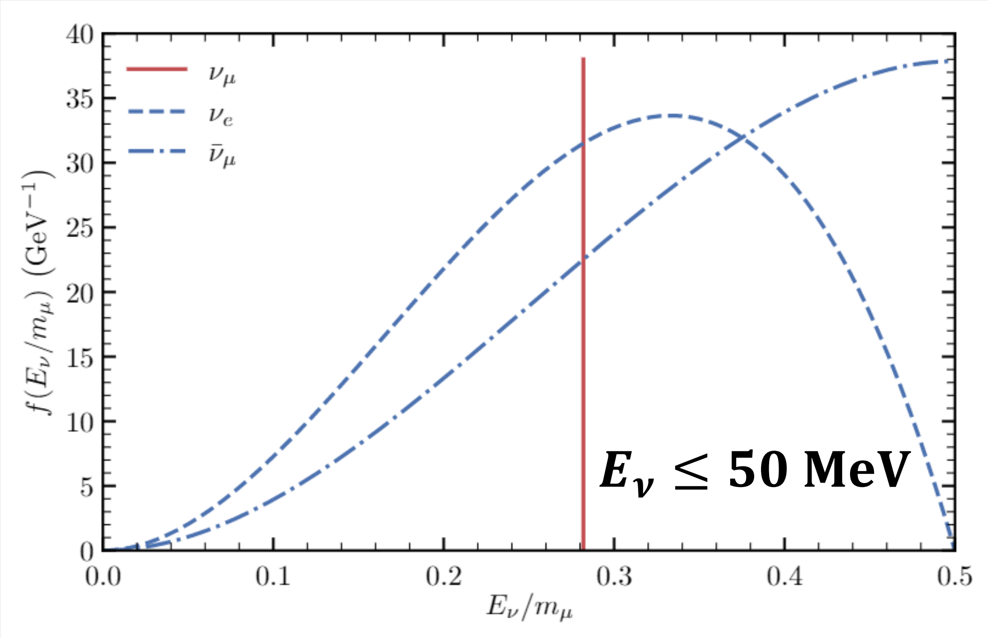






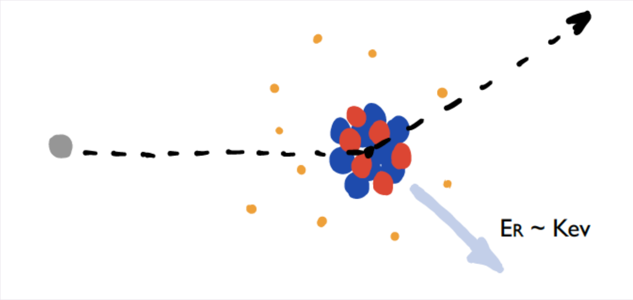
[Miranda et al. 2008.02759 (2020)]

Experiment	Mass [ton]	E_{th} [keV$_{\rm nr}$]	NPOT [10^{23}/yr]	r	L [m]	$\sigma_{\rm sys}$
CENNS610	0.61	~ 20	1.5	0.08	28.4	8.5%
ESS10	0.01	0.1	2.8	0.3	20	5%
CCM	7	10	0.177	0.0425	20	5%
ESS	1	20	2.8	0.3	20	5%

not low energy thresholds

CEνNS at Direct Detection

$$\frac{\mathrm{d}R}{\mathrm{d}E_R} = n_T \sum_{\nu_\alpha} \int_{E_\nu^{\min}} \frac{\mathrm{d}\phi_{\nu_e}}{\mathrm{d}E_\nu} \, P(\nu_e \to \nu_\alpha) \, \frac{\mathrm{d}\sigma_{\nu_\alpha T}}{\mathrm{d}E_R} \, \mathrm{d}E_\nu$$



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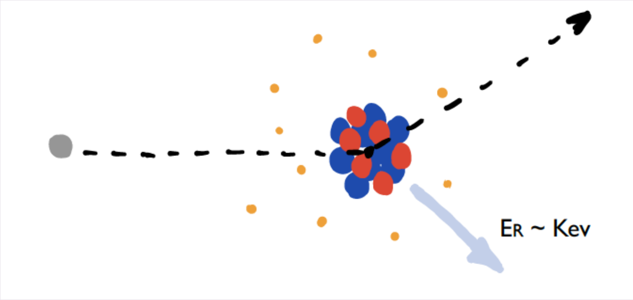

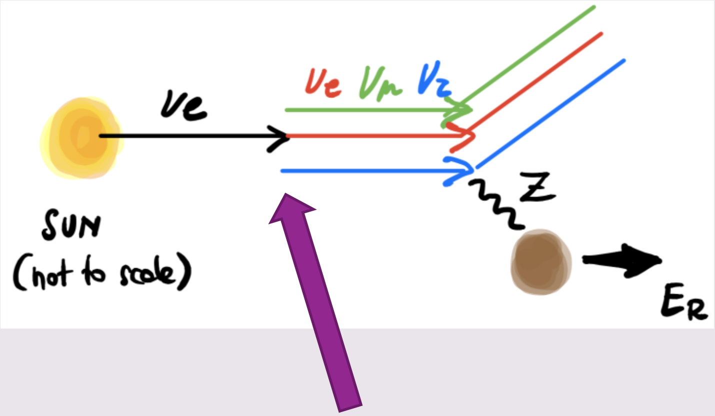

Neutrino oscillations

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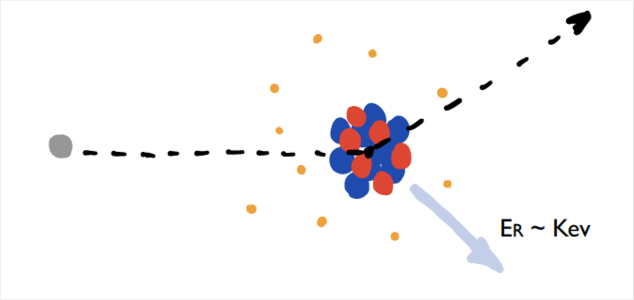

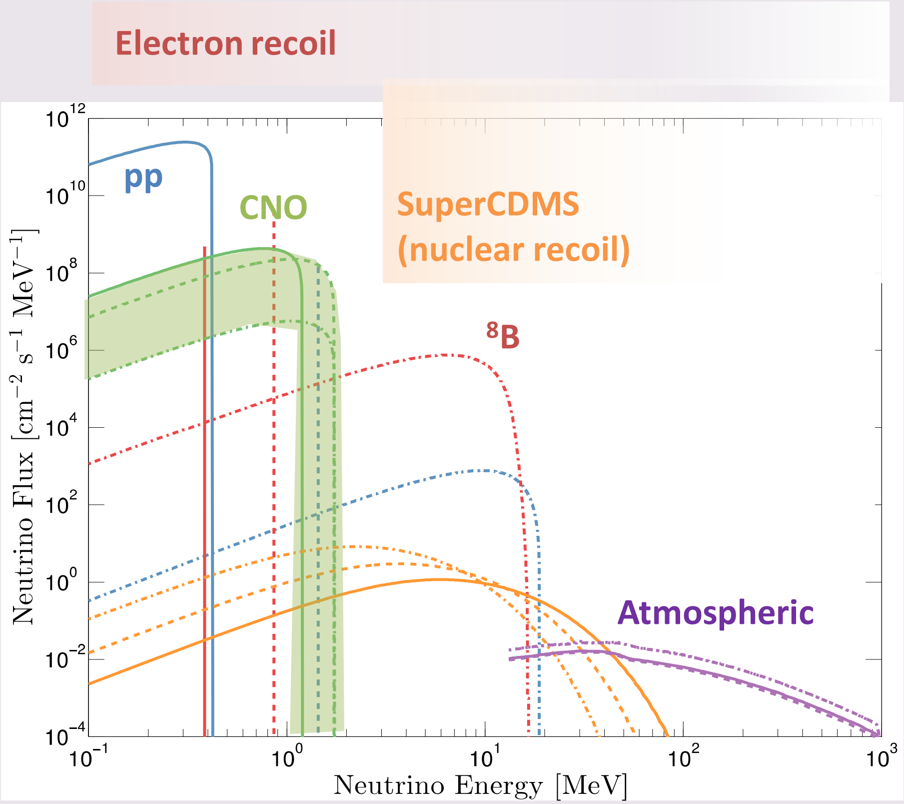

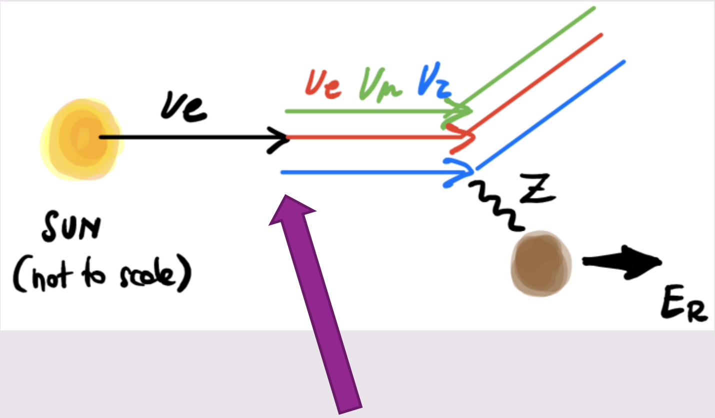

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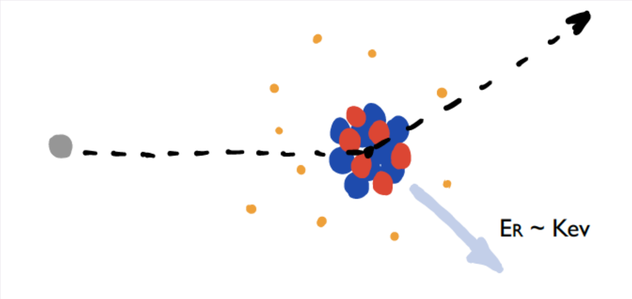

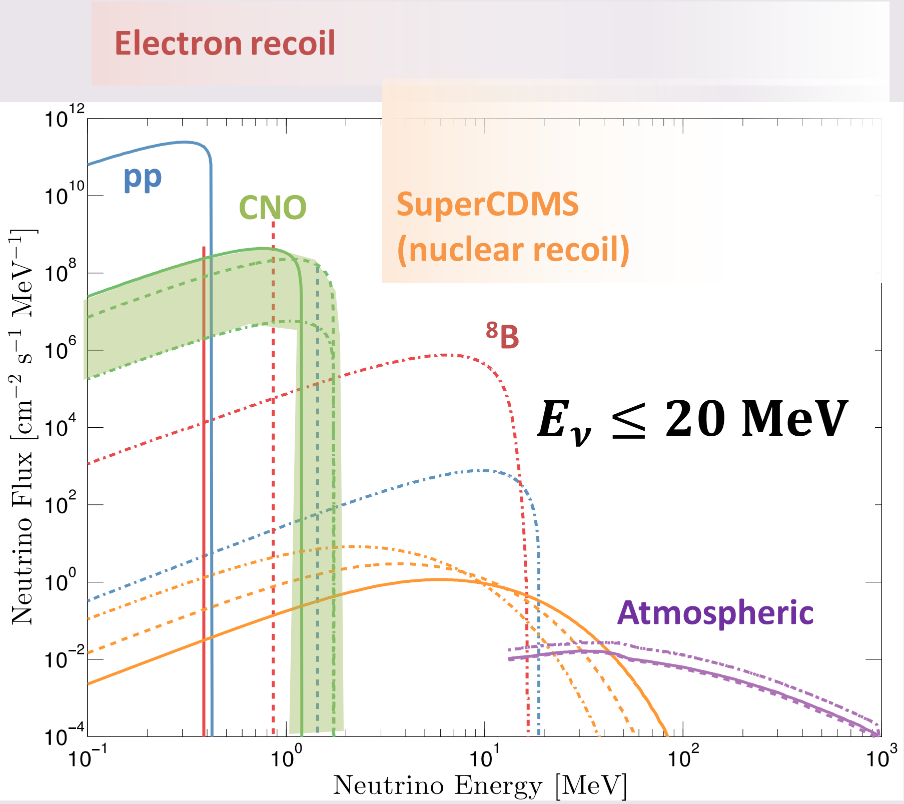

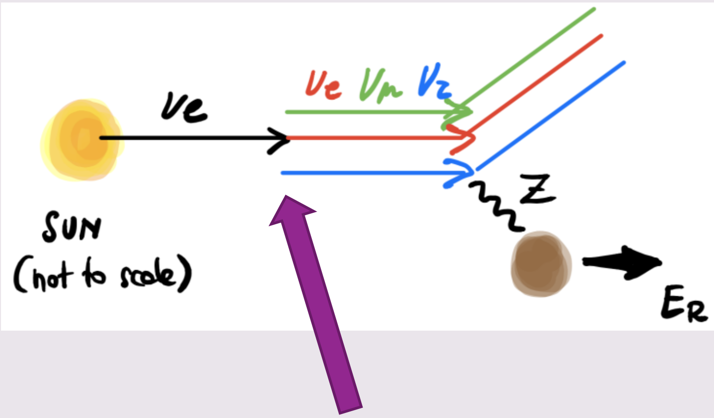

Neutrino oscillations

Different experimental setups…

Spallation Sources	Direct Detection

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…so why not **combine** them?

Sterile Baryonic Neutrino (SBN)

[Pospelov 1103.3261 (2011)]

$$L_{SBN} \supset g_{Z'} \frac{1}{3} \sum_{q} \bar{q}\, \gamma_\mu Z'^{\mu} q + g_{Z'} \overline{\nu_b} \gamma_\mu Z'^{\mu} \nu_b$$

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ν_b: sterile baryonic neutrino (m_4)

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PARAMETER SPACE

$\boldsymbol{g_{Z'}}$, $\boldsymbol{m_{Z'}}$

$\boldsymbol{m_4}$, $\boldsymbol{|U_{e4}|}$, $\boldsymbol{|U_{\mu 4}|}$, $\boldsymbol{|U_{\tau 4}|}$

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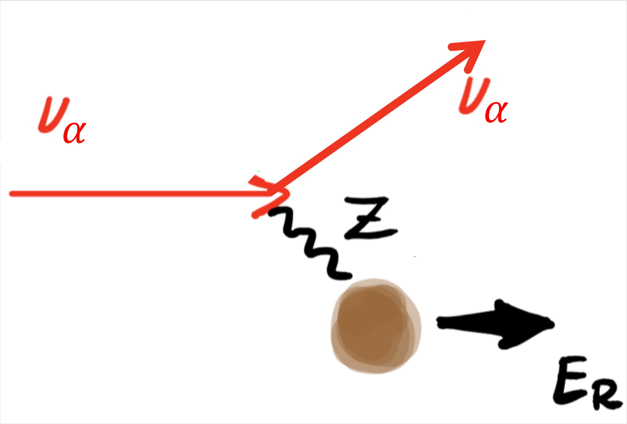

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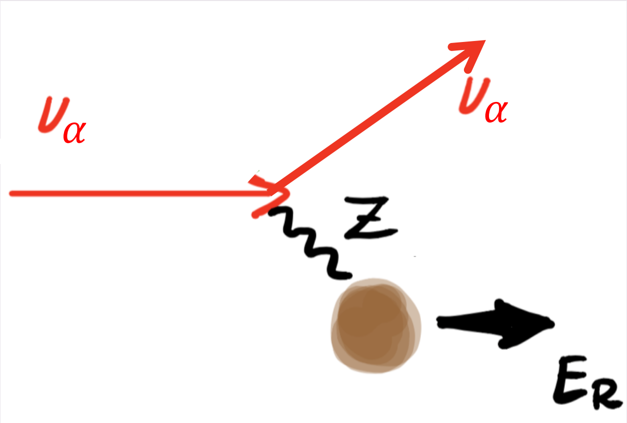

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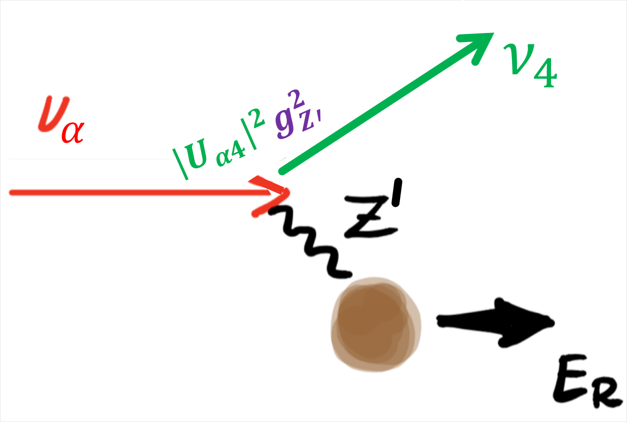

SBN predicts **upscattering** process:

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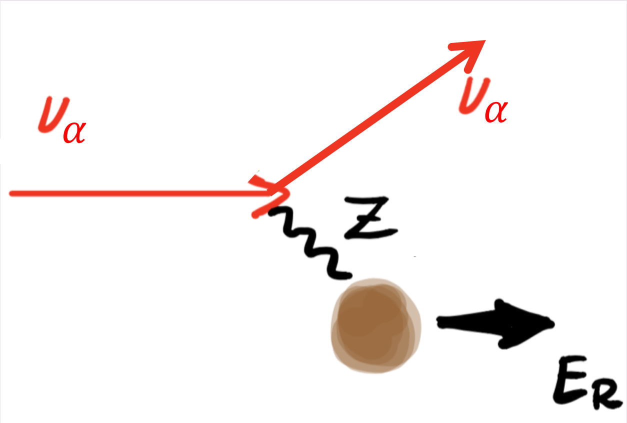

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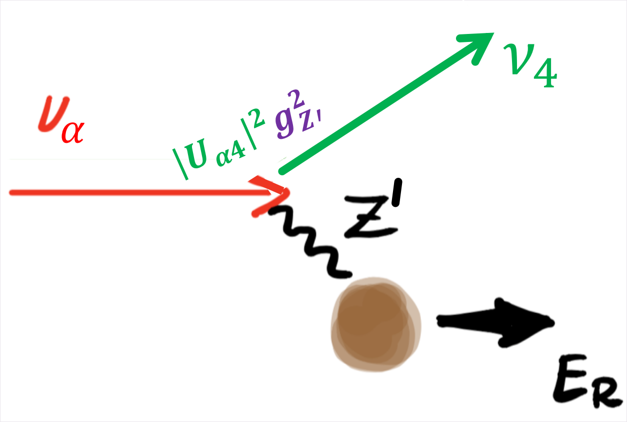

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that may **modify** the expected **recoil spectrum!**

Let's fix **our scenario…**

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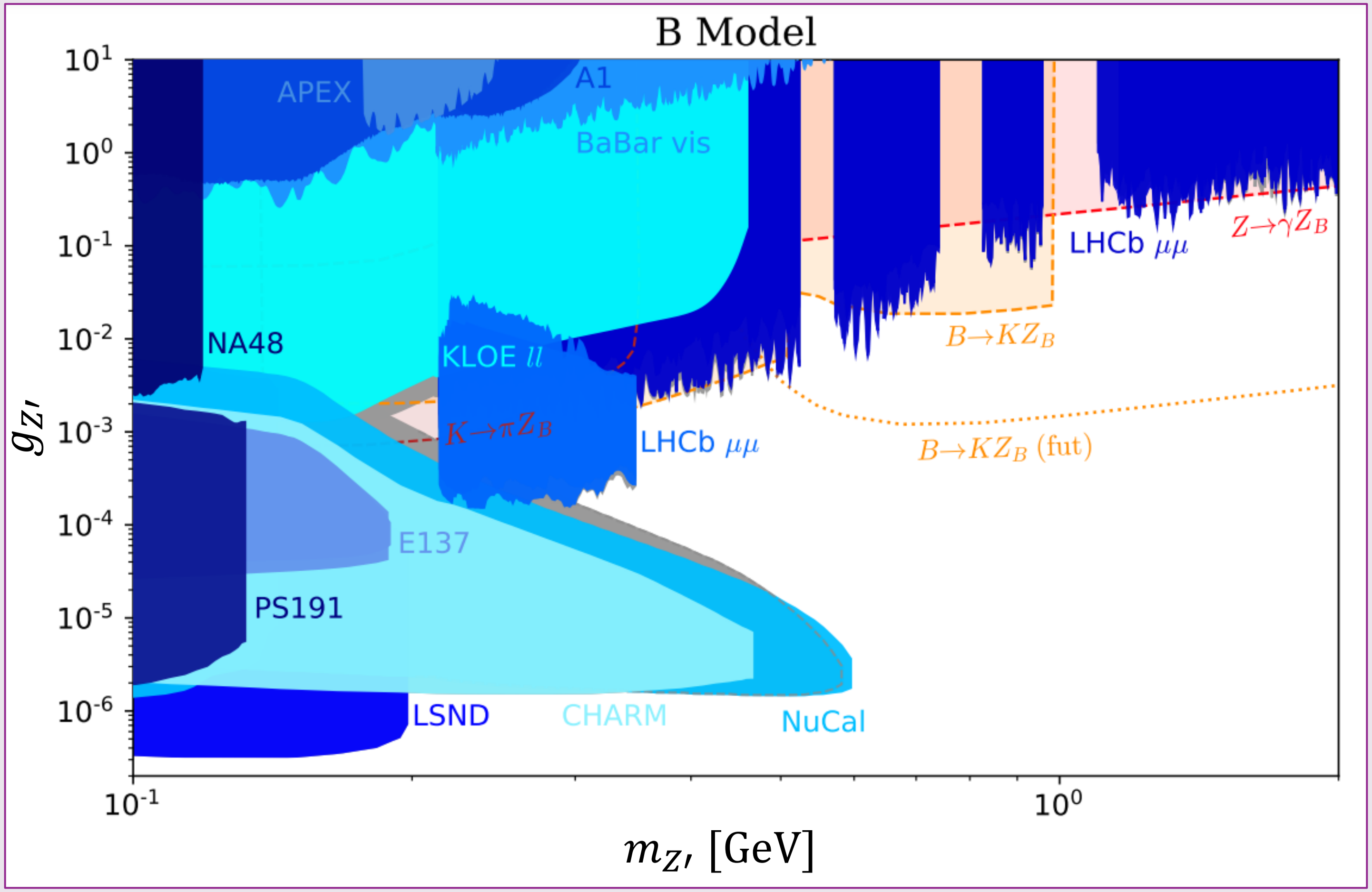

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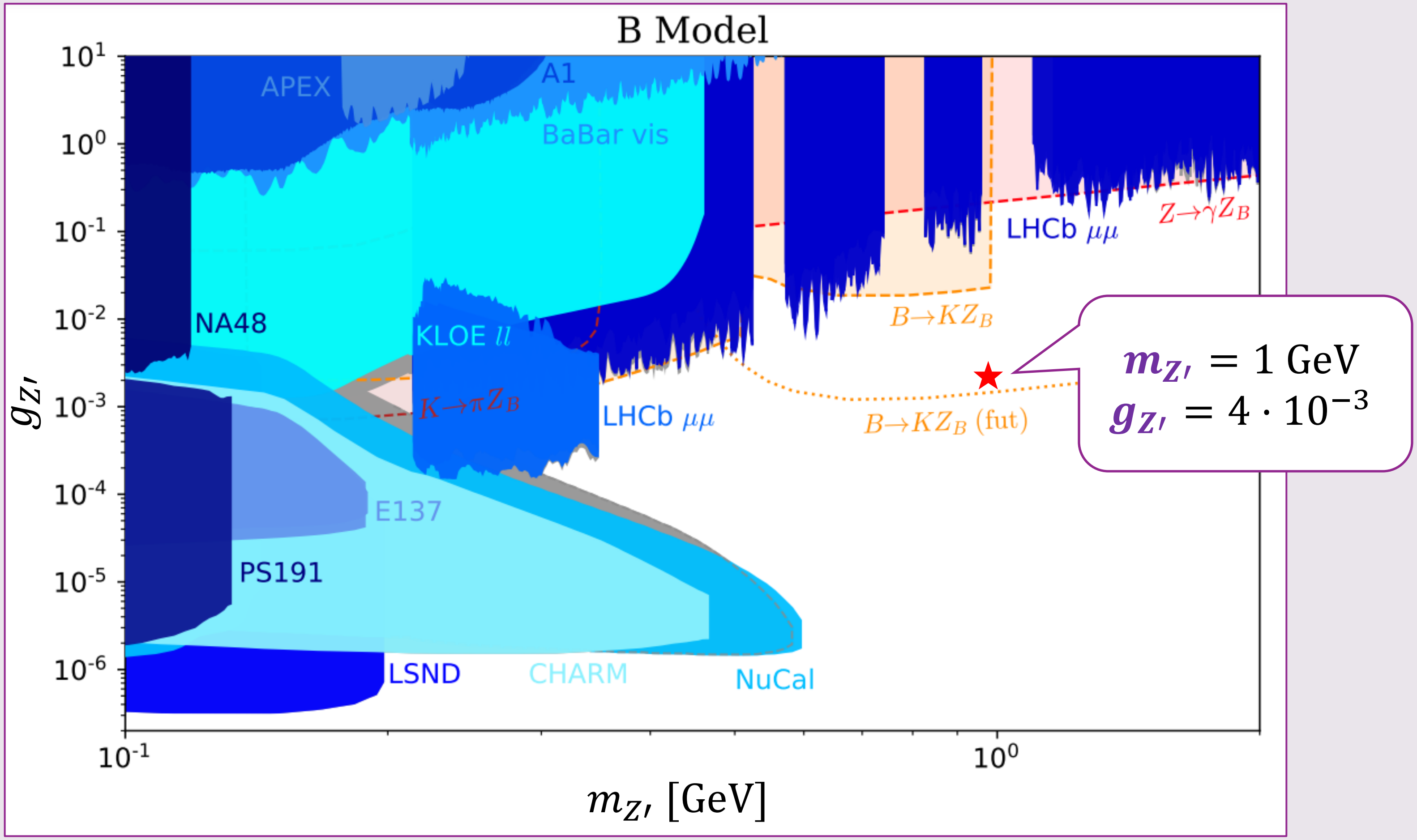

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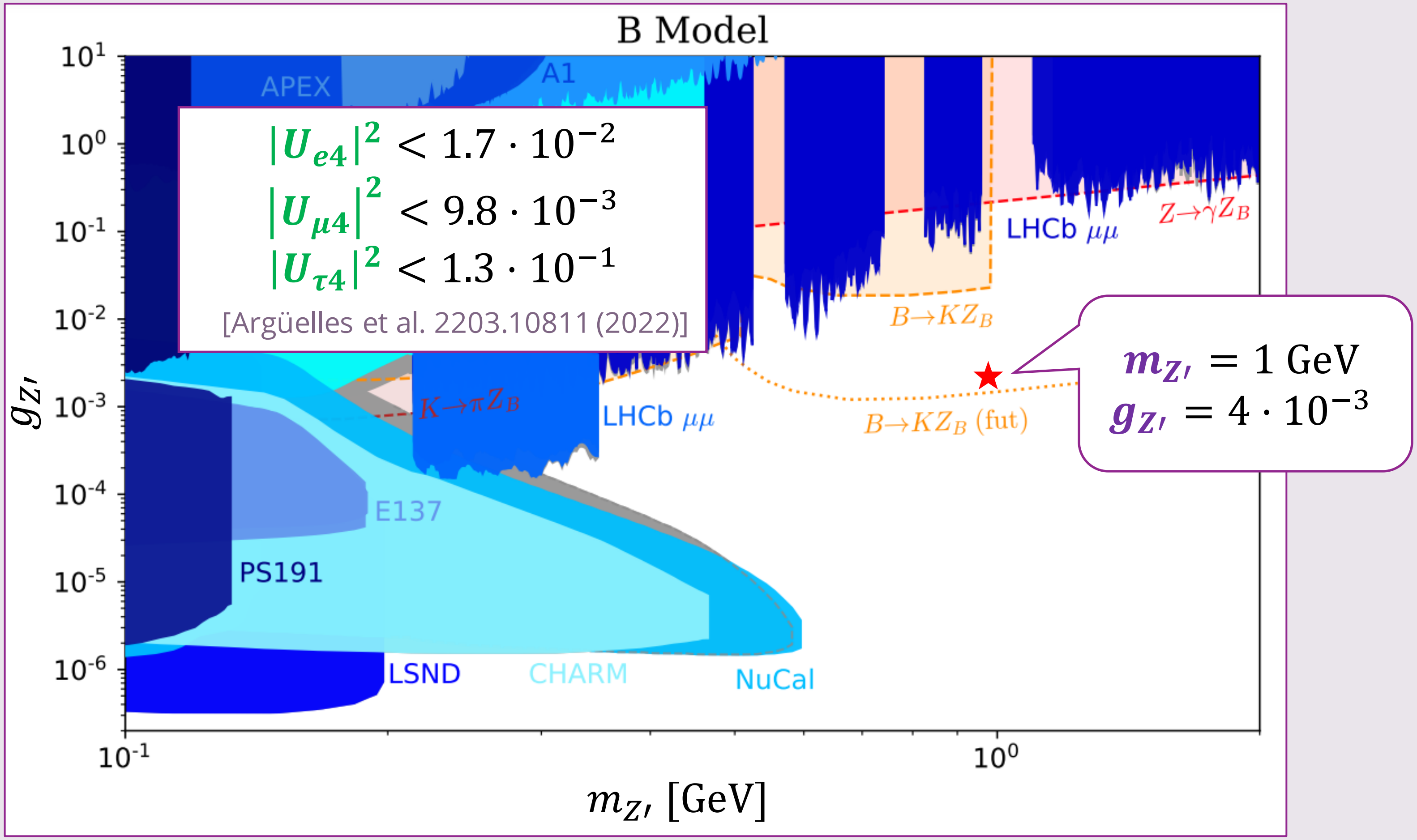

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Let’s fix some **benchmark points...**

	m_4 [GeV]	$\|U_{e4}\|^2$	$\|U_{\mu 4}\|^2$	$\|U_{\tau 4}\|^2$
BP1a	2×10^{-3}	0	9×10^{-3}	0
BP1d	2×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP2a	9×10^{-3}	0	9×10^{-3}	0
BP2b	9×10^{-3}	0	9×10^{-3}	9×10^{-4}
BP2c	9×10^{-3}	0	9×10^{-3}	4×10^{-3}
BP2d	9×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP3a	20×10^{-3}	0	9×10^{-3}	0
BP4a	40×10^{-3}	0	9×10^{-3}	0
BP5a	60×10^{-3}	0	9×10^{-3}	0

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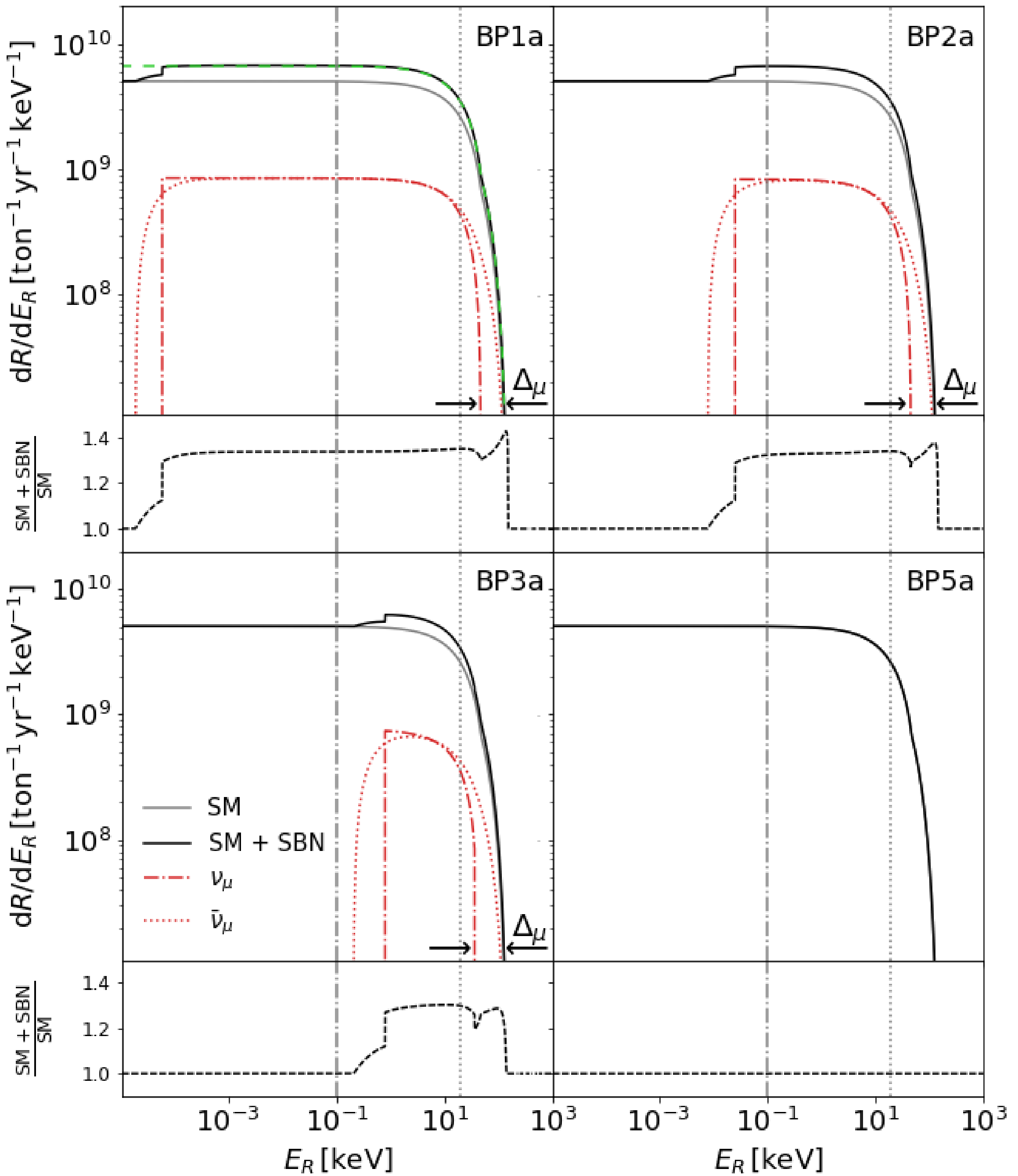
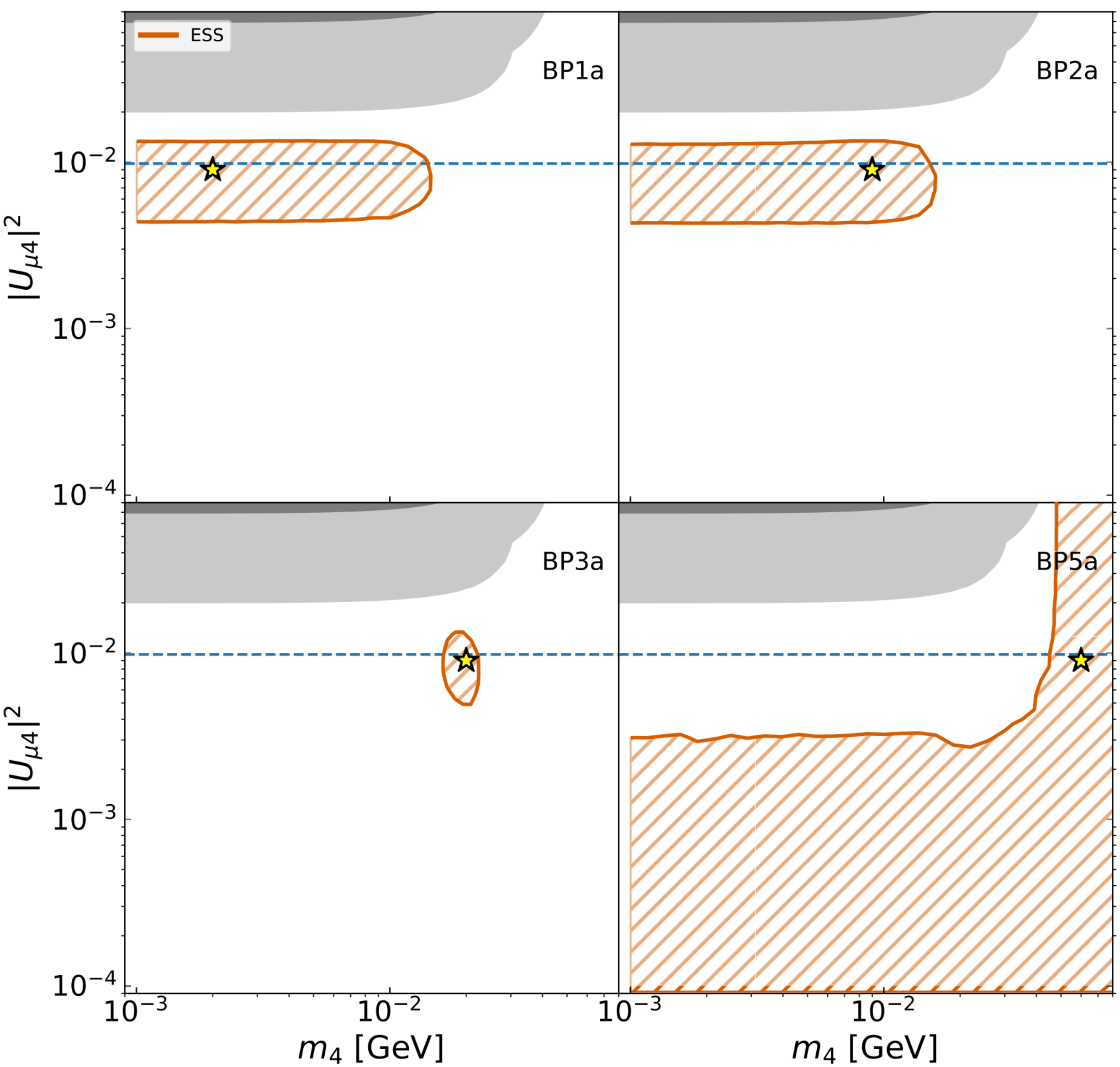
BP2c	9×10^{-3}	0	9×10^{-3}	4×10^{-3}
BP2d	9×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP3a	20×10^{-3}	0	9×10^{-3}	0
BP4a	40×10^{-3}	0	9×10^{-3}	0
BP5a	60×10^{-3}	0	9×10^{-3}	0

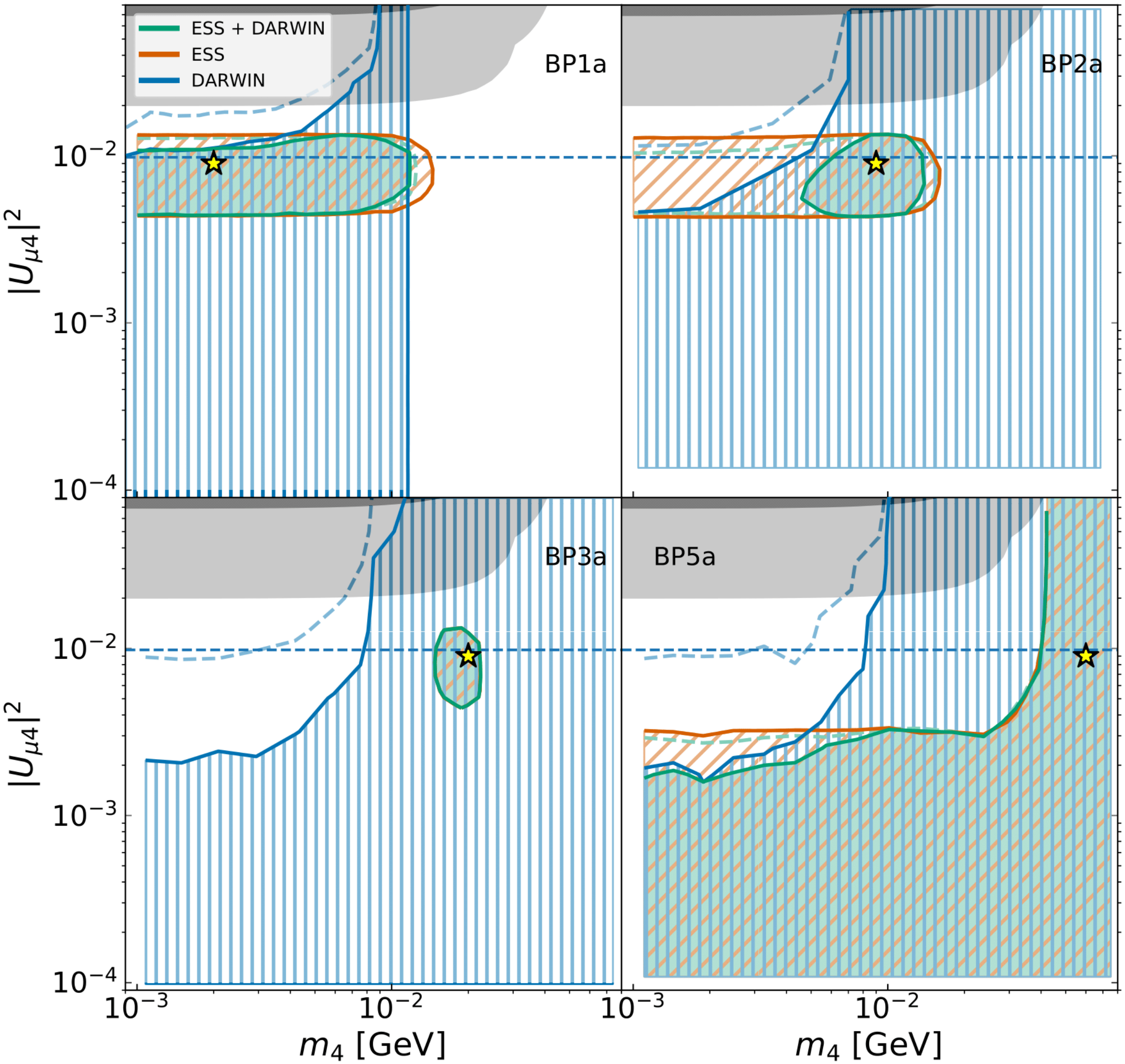
Predicted SS energy spectra

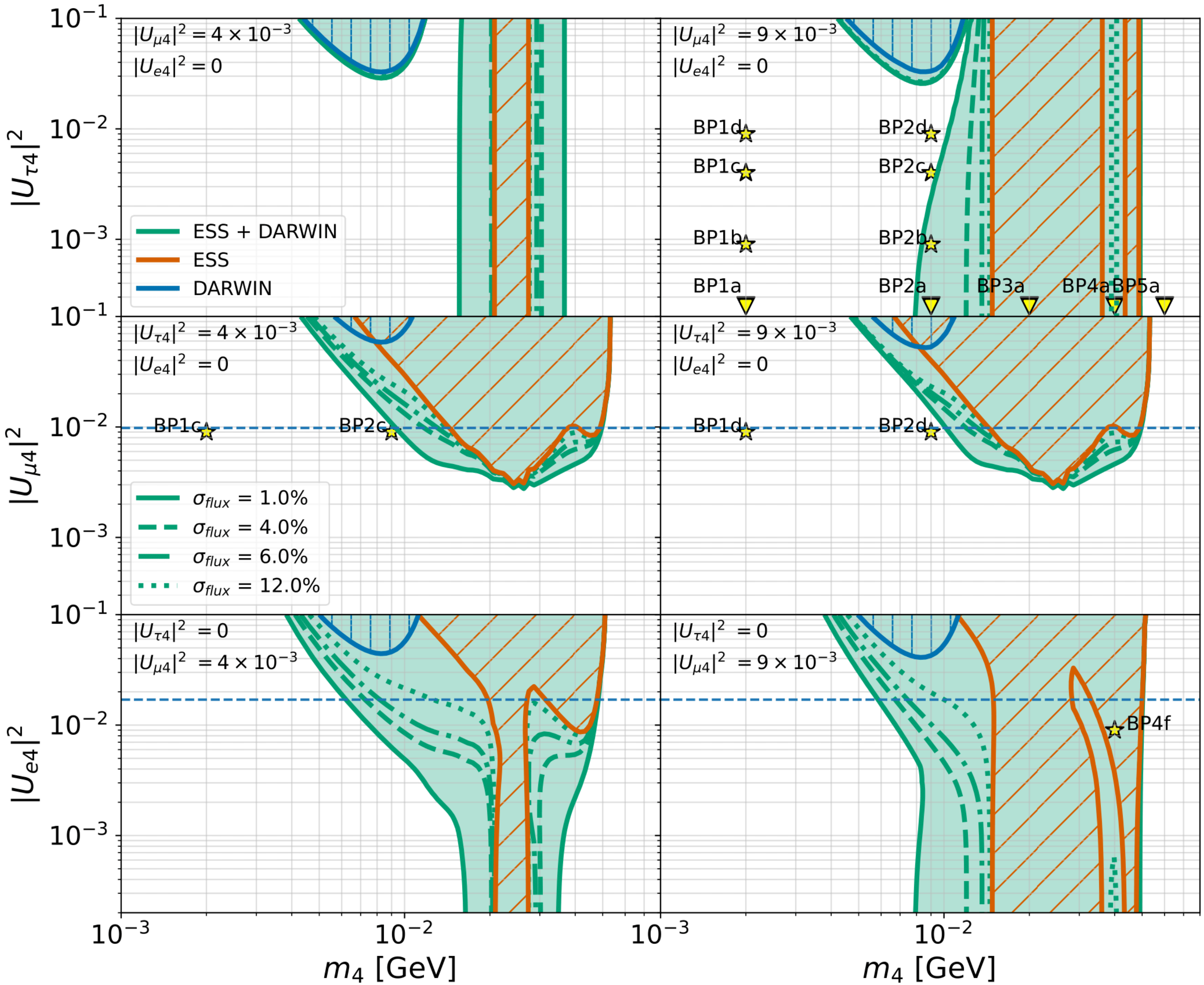


SS Profile likelihood results



SS+DD Profile likelihood results





$|U_{\mu 4}|^2 = 4 \times 10^{-3}$
$|U_{e4}|^2 = 0$
$|U_{\mu 4}|^2 = 9 \times 10^{-3}$
$|U_{e4}|^2 = 0$
$|U_{\tau 4}|^2 = 4 \times 10^{-3}$
$|U_{e4}|^2 = 0$
$|U_{\tau 4}|^2 = 9 \times 10^{-3}$
$|U_{e4}|^2 = 0$
$|U_{\tau 4}|^2 = 0$
$|U_{\mu 4}|^2 = 4 \times 10^{-3}$
$|U_{\tau 4}|^2 = 0$
$|U_{\mu 4}|^2 = 9 \times 10^{-3}$
ESS + DARWIN
ESS
DARWIN
$\sigma_{flux} = 1.0\%$
$\sigma_{flux} = 4.0\%$
$\sigma_{flux} = 6.0\%$
$\sigma_{flux} = 12.0\%$
BP1a
BP1b
BP1c
BP1d
BP2a
BP2b
BP2c
BP2d
BP3a
BP4a
BP5a
BP4f
$|U_{\tau 4}|^2$
$|U_{\mu 4}|^2$
$|U_{e4}|^2$
m_4 [GeV]

Conclusions

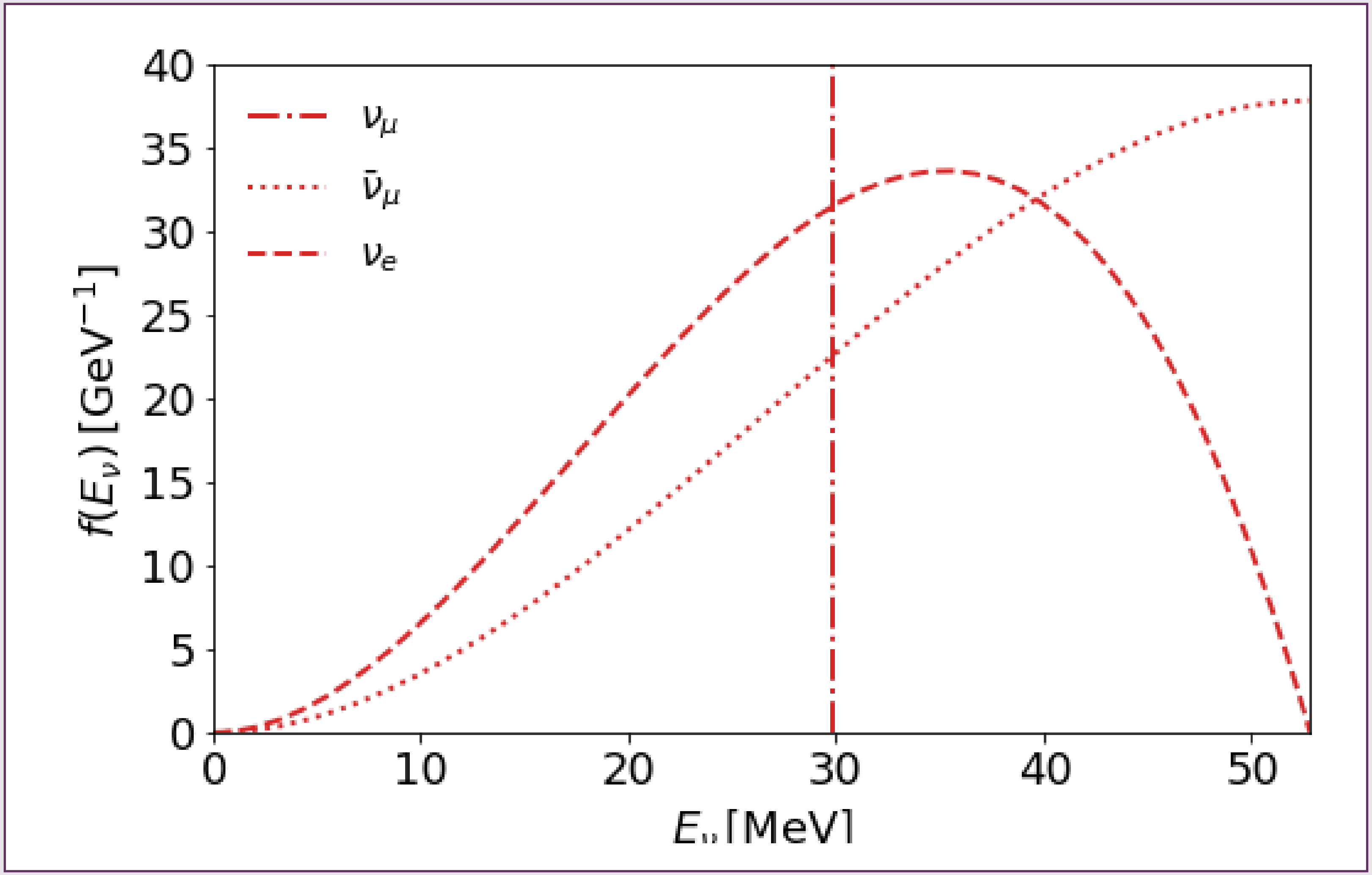
- Sterile neutrino models can be **probed** with Spallation Source (**SS**) and Direct Detection (**DD**) experiments.
- **DD** will be able to access to very **low recoil energies, all the neutrino flavours** but not big masses.
- **SS** will be able to access to **heavier sterile neutrinos** but not to all neutrino flavours.
- **Combining** DD and SS may help…
  - **improving the significance,**
  - constraining the parameter space and allowing **parameter reconstruction (specially in the neutrino mass m_4** and in τ mixing),
  - and allowing model discrimination (Sterile Baryonic Neutrino vs NSI).
- DD must reach smaller thresholds and the **uncertainty in solar neutrino fluxes need to be reduced** in order to be competitive.

Thanks for your attention!

Backup Slides

Direct Detection and
Spallation Experiments to test the
Baryonic Sterile Neutrino

SS experiment **fluxes**

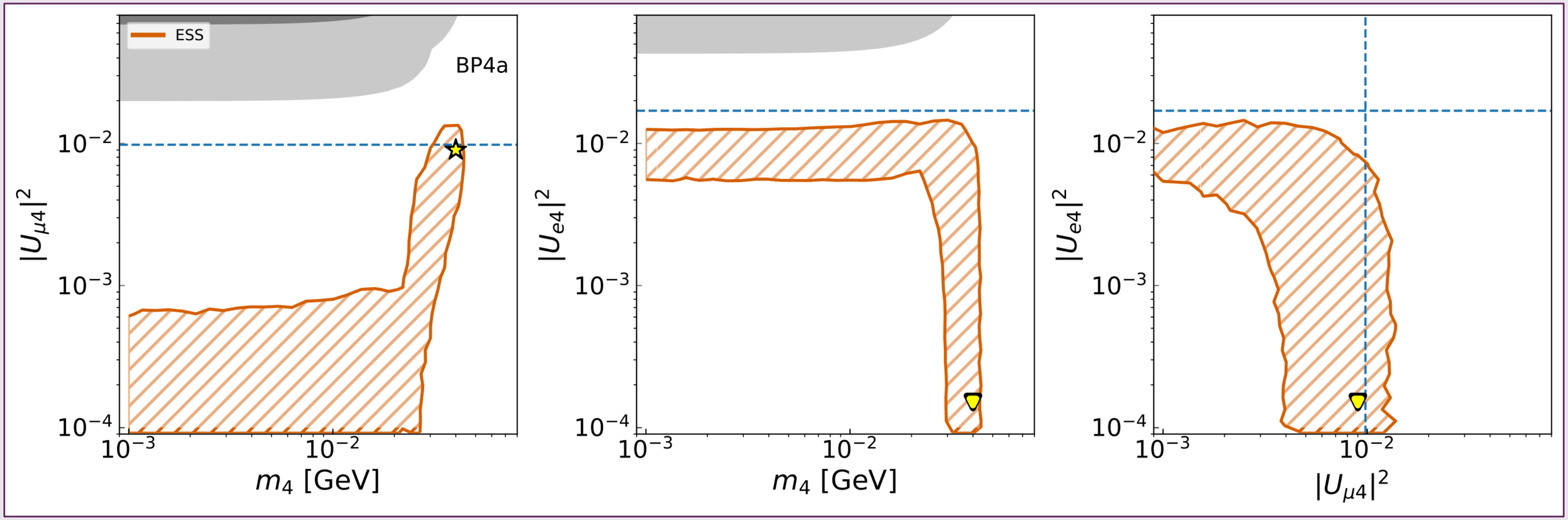


DD + SS: why it's a good idea

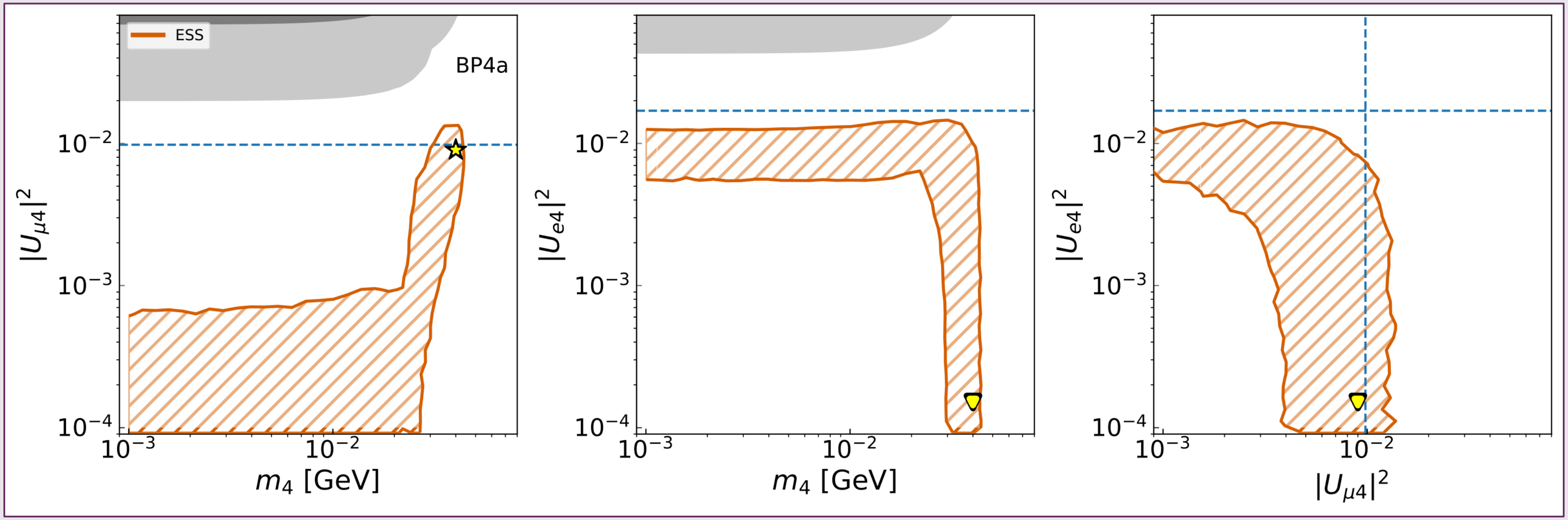
1. Increase the **statistical significance** of a prospective discovery.
2. Improve the **parameter reconstruction** of the model.
3. Allow to **discriminate** between our model and other **models** that can give similar experimental evidence.

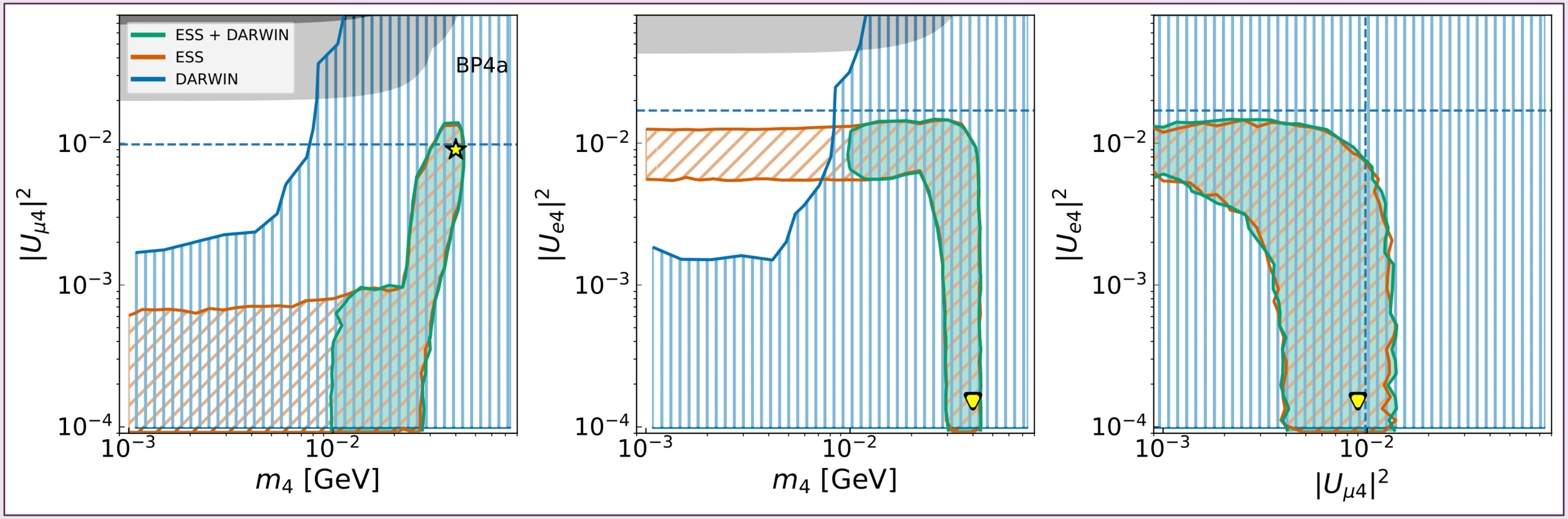
Let's fix some **benchmark points...**

	m_4 [GeV]	$\|U_{e4}\|^2$	$\|U_{\mu 4}\|^2$	$\|U_{\tau 4}\|^2$
BP1a	2×10^{-3}	0	9×10^{-3}	0
BP1d	2×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP2a	9×10^{-3}	0	9×10^{-3}	0
BP2b	9×10^{-3}	0	9×10^{-3}	9×10^{-4}
BP2c	9×10^{-3}	0	9×10^{-3}	4×10^{-3}
BP2d	9×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP3a	20×10^{-3}	0	9×10^{-3}	0
BP4a	40×10^{-3}	0	9×10^{-3}	0
BP5a	60×10^{-3}	0	9×10^{-3}	0



ESS
BP4a
$|U_{\mu4}|^2$
m_4 [GeV]
$|U_{e4}|^2$
m_4 [GeV]
$|U_{e4}|^2$
$|U_{\mu4}|^2$
10^{-2}
10^{-3}
10^{-4}

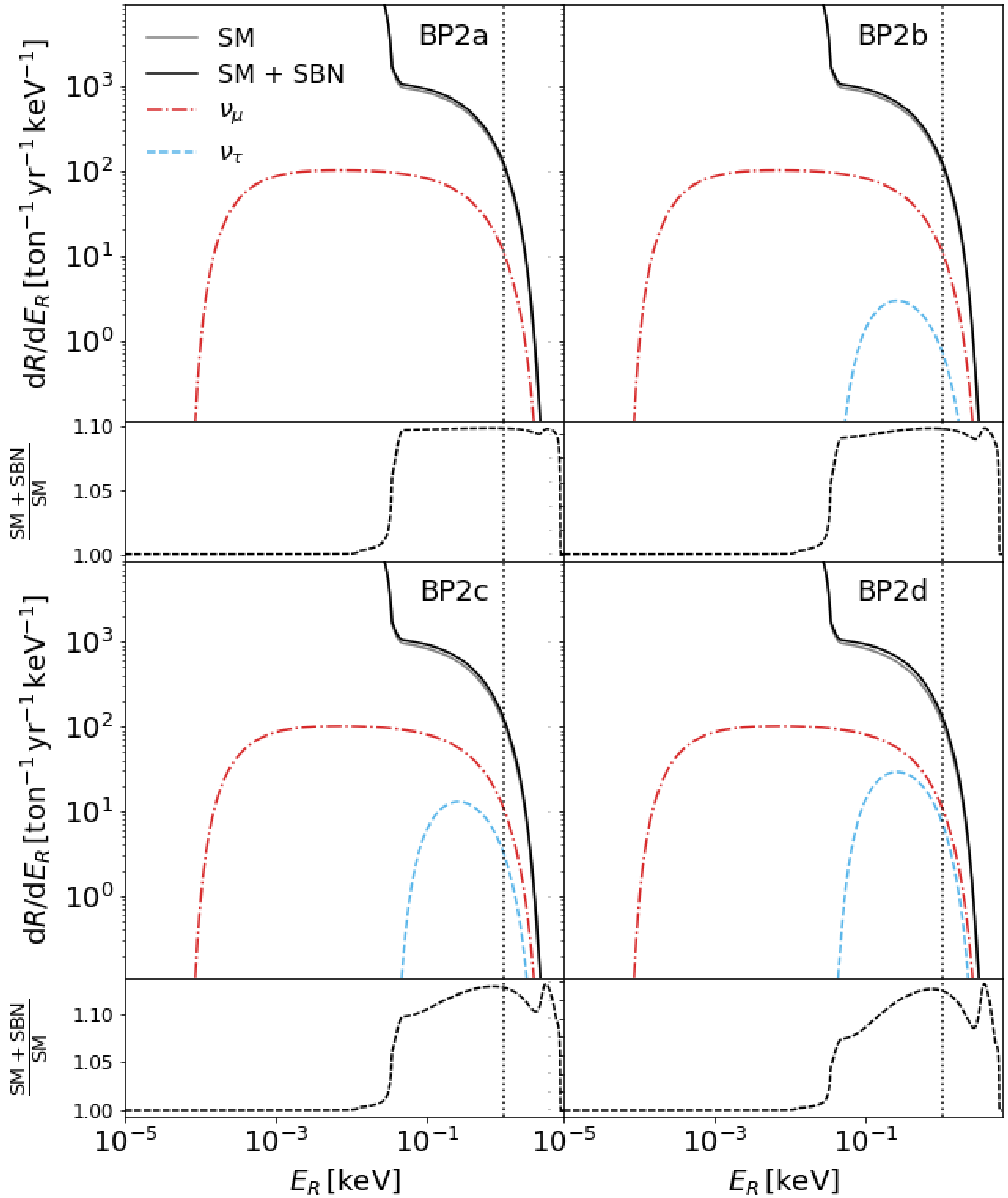

ESS
BP4a
10^{-2}
10^{-3}
10^{-4}
$|U_{\mu4}|^2$
$|U_{e4}|^2$
m_4 [GeV]
$|U_{\mu4}|^2$


ESS + DARWIN
ESS
DARWIN
BP4a
$|U_{\mu4}|^2$
$|U_{e4}|^2$
m_4 [GeV]
$|U_{\mu4}|^2$

Let's fix some **benchmark points...**

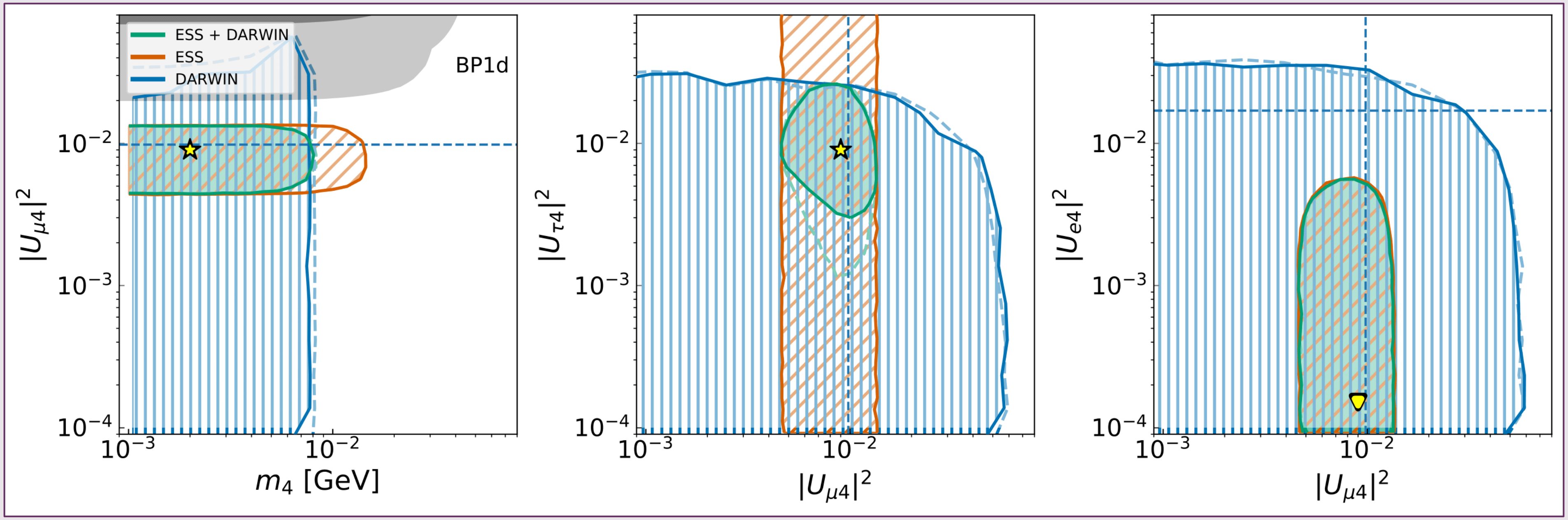
	m_4 [GeV]	$\lvert U_{e4}\rvert^2$	$\lvert U_{\mu 4}\rvert^2$	$\lvert U_{\tau 4}\rvert^2$
BP1a	2×10^{-3}	0	9×10^{-3}	0
BP1d	2×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP2a	9×10^{-3}	0	9×10^{-3}	0
BP2b	9×10^{-3}	0	9×10^{-3}	9×10^{-4}
BP2c	9×10^{-3}	0	9×10^{-3}	4×10^{-3}
BP2d	9×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP3a	20×10^{-3}	0	9×10^{-3}	0
BP4a	40×10^{-3}	0	9×10^{-3}	0
BP5a	60×10^{-3}	0	9×10^{-3}	0

Predicted DD energy spectra



Let's fix some **benchmark points...**

	m_4 [GeV]	$\|U_{e4}\|^2$	$\|U_{\mu 4}\|^2$	$\|U_{\tau 4}\|^2$
BP1a	2×10^{-3}	0	9×10^{-3}	0
BP1d	2×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP2a	9×10^{-3}	0	9×10^{-3}	0
BP2b	9×10^{-3}	0	9×10^{-3}	9×10^{-4}
BP2c	9×10^{-3}	0	9×10^{-3}	4×10^{-3}
BP2d	9×10^{-3}	0	9×10^{-3}	9×10^{-3}
BP3a	20×10^{-3}	0	9×10^{-3}	0
BP4a	40×10^{-3}	0	9×10^{-3}	0
BP5a	60×10^{-3}	0	9×10^{-3}	0


ESS + DARWIN
ESS
DARWIN
BP1d
$|U_{\mu 4}|^2$
m_4 [GeV]
$|U_{\tau 4}|^2$
$|U_{\mu 4}|^2$
$|U_{e4}|^2$
$|U_{\mu 4}|^2$
10^{-2}
10^{-3}
10^{-4}